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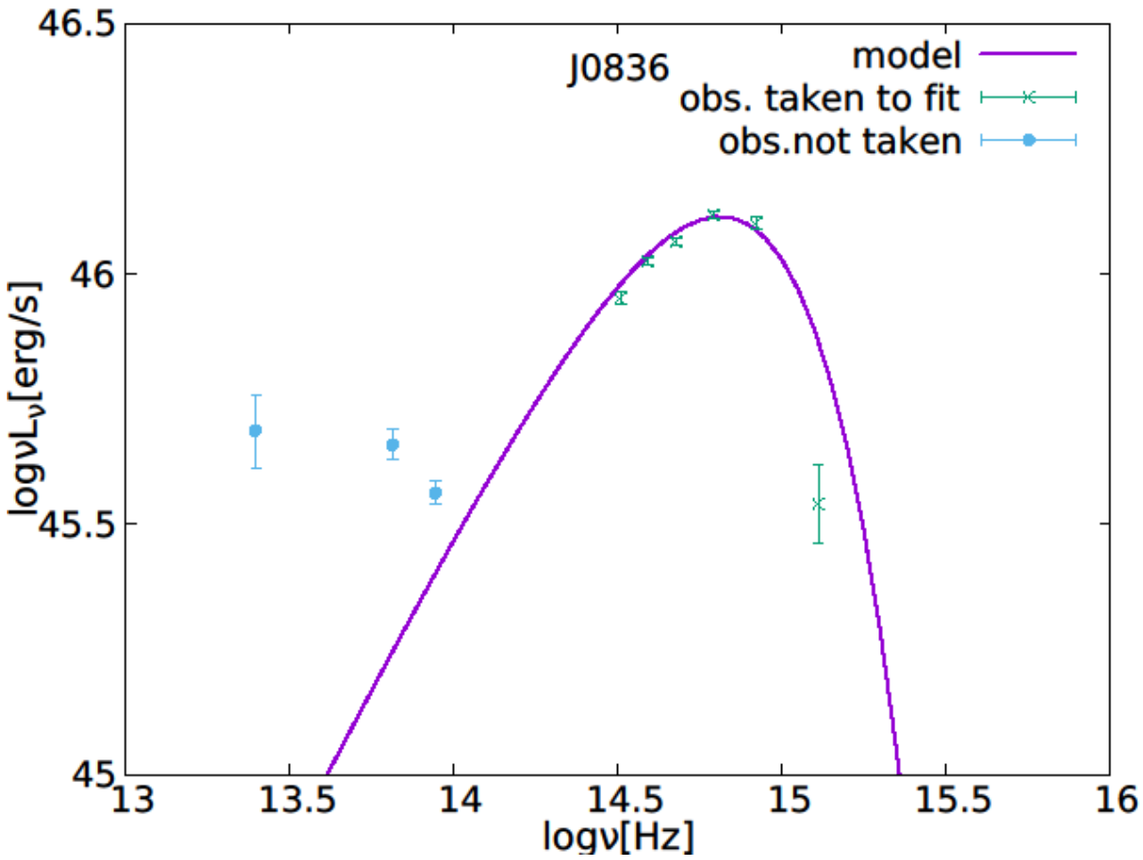
We study weak emission-line quasars (WLQ). The primary goal of this work is to determine four parameters (mass of black hole, accretion rate, spin and inclination) by the geometrically thin and optically thick Novikov-Thorne (NT) model. In this work, we use the numerical code based on the NT equations. The model of continuum of accretion disk is placed on the Shakura-Sunayaev model but instead of a=0 they used non-zero spin.



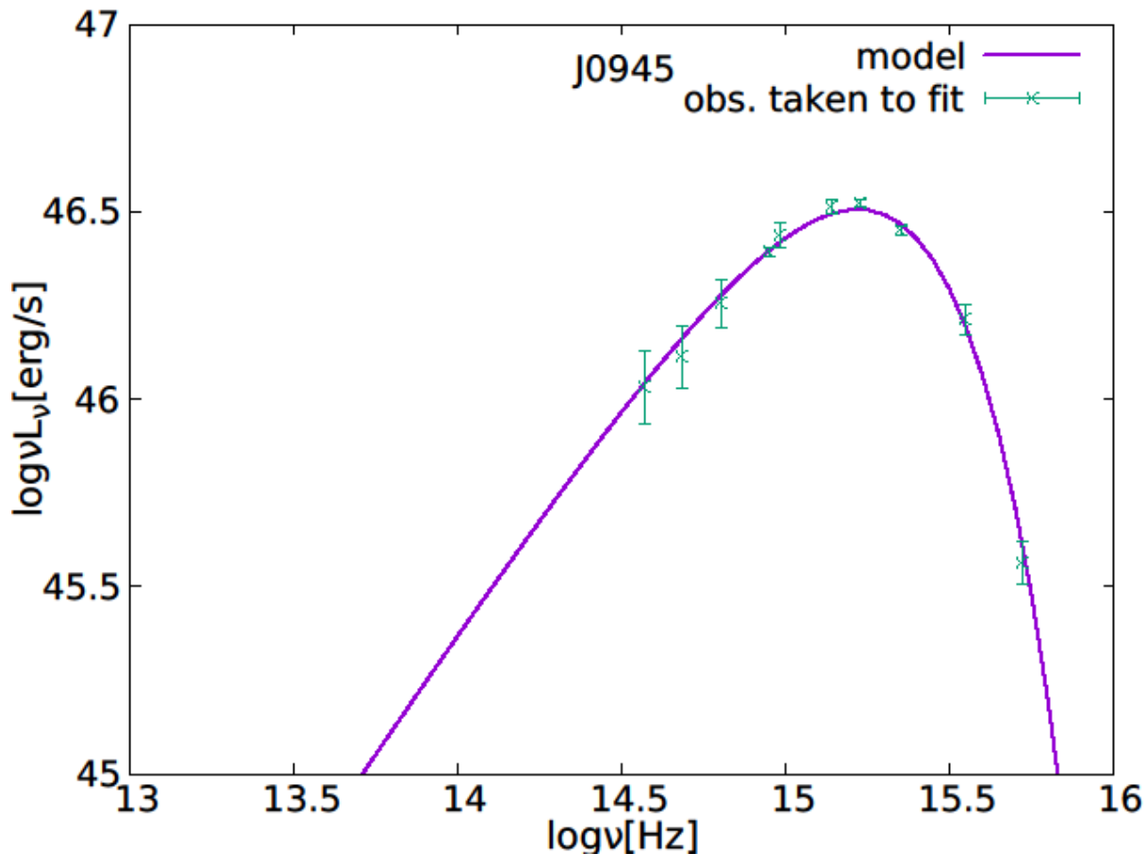


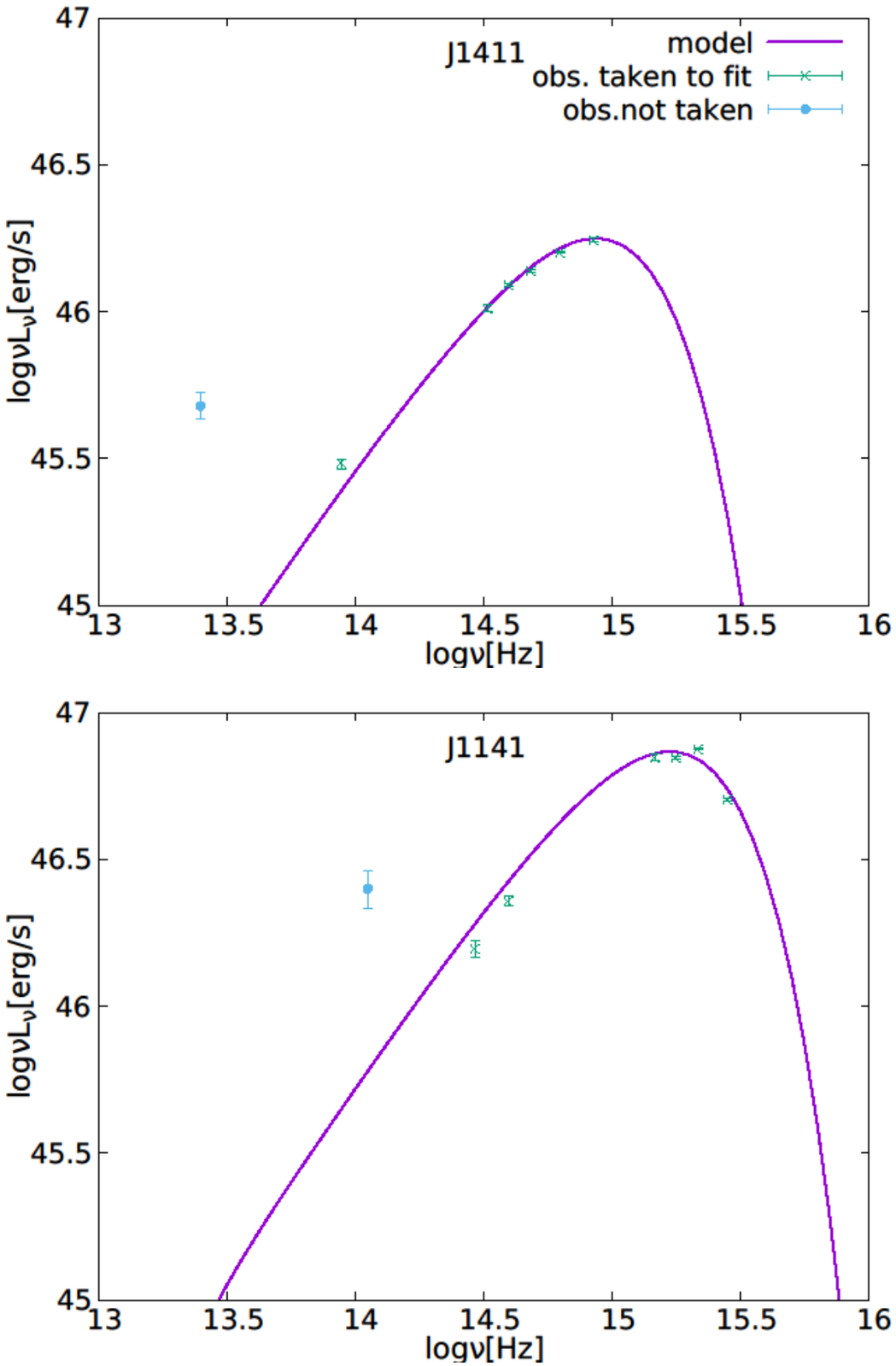
Table : Table of best parameters

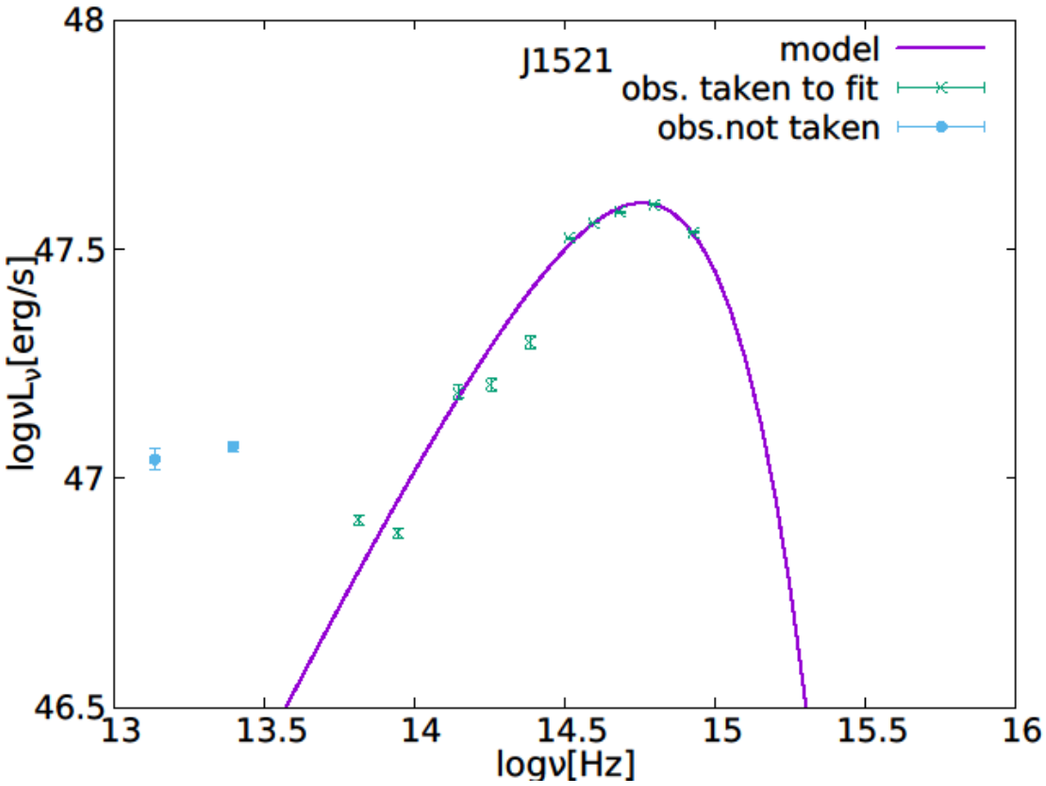
Name	M_{BH}	$\dot{m}$	χ^2/d.o.f	Bayesian M_{BH}
J0836	$(1.00^{+1.64}_{0.30})E+10$	0.035 ± 0.034	4.42	$(2.00 \pm 1.09)E+10$
J0945	$(6.30 \pm 1.97)E+9$	$0.146^{+0.020}_{-0.020}$	1.19	$(6.30 \pm 3.96)E+9$
J1141	$(3.80 \pm 2.76)E+9$	$0.420^{+0.399}_{-0.151}$	2.99	$(4.00^{+8.18}_{-3.00})E+9$
J1411	$(7.85 \pm 2.20)E+9$	$0.058^{+0.027}_{-0.021}$	1.131	$(7.90 \pm 2.02)E+9$
J1521	$(7.90^{+2.55}_{-0.85})E+10$	0.114 ± 0.013	1.61	$(1.30 \pm 0.64)E+11$

Models of continuum of weak emission-line quasars

The main conclusions are:

- We are able to describe mass of black hole and accretion rate of WLQs,
- We compared parameters of WLQs by χ^2 and Bayesian method





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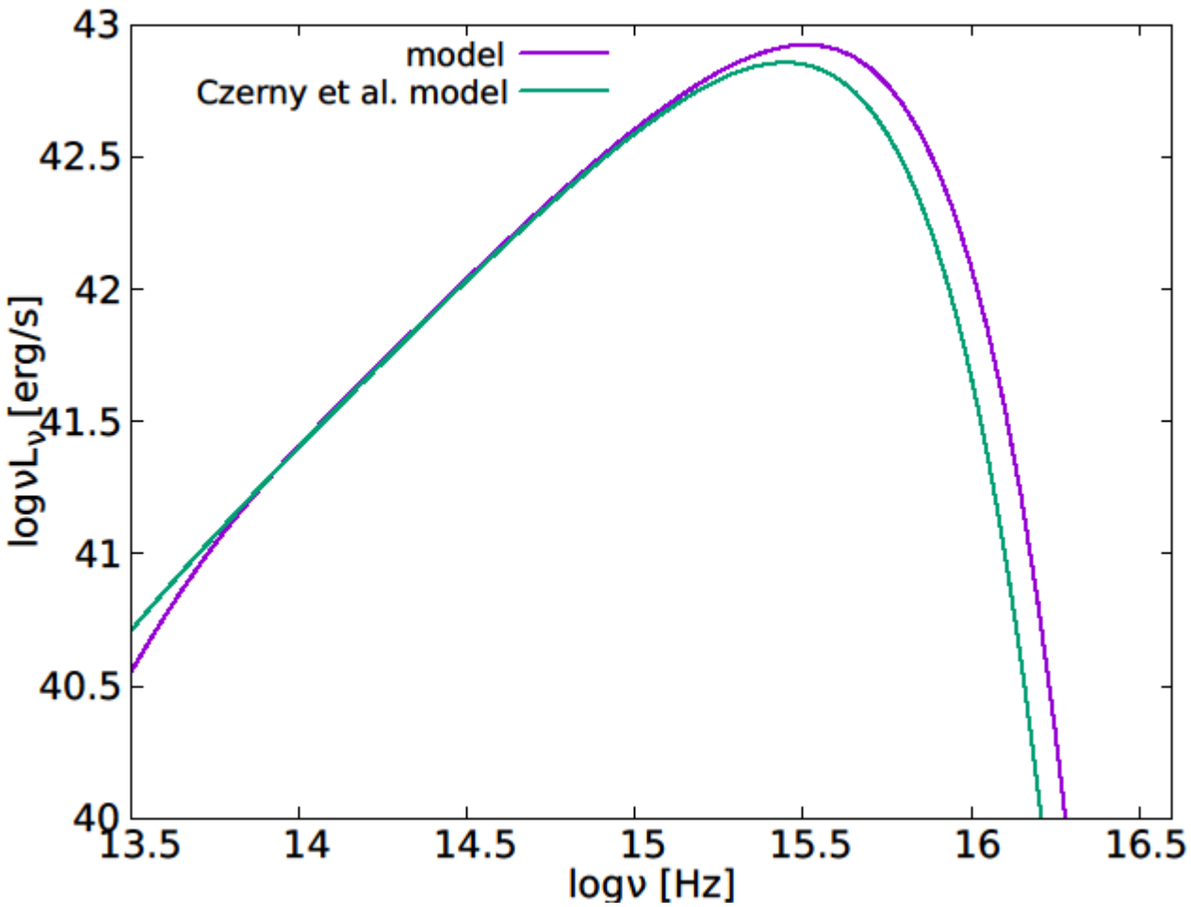
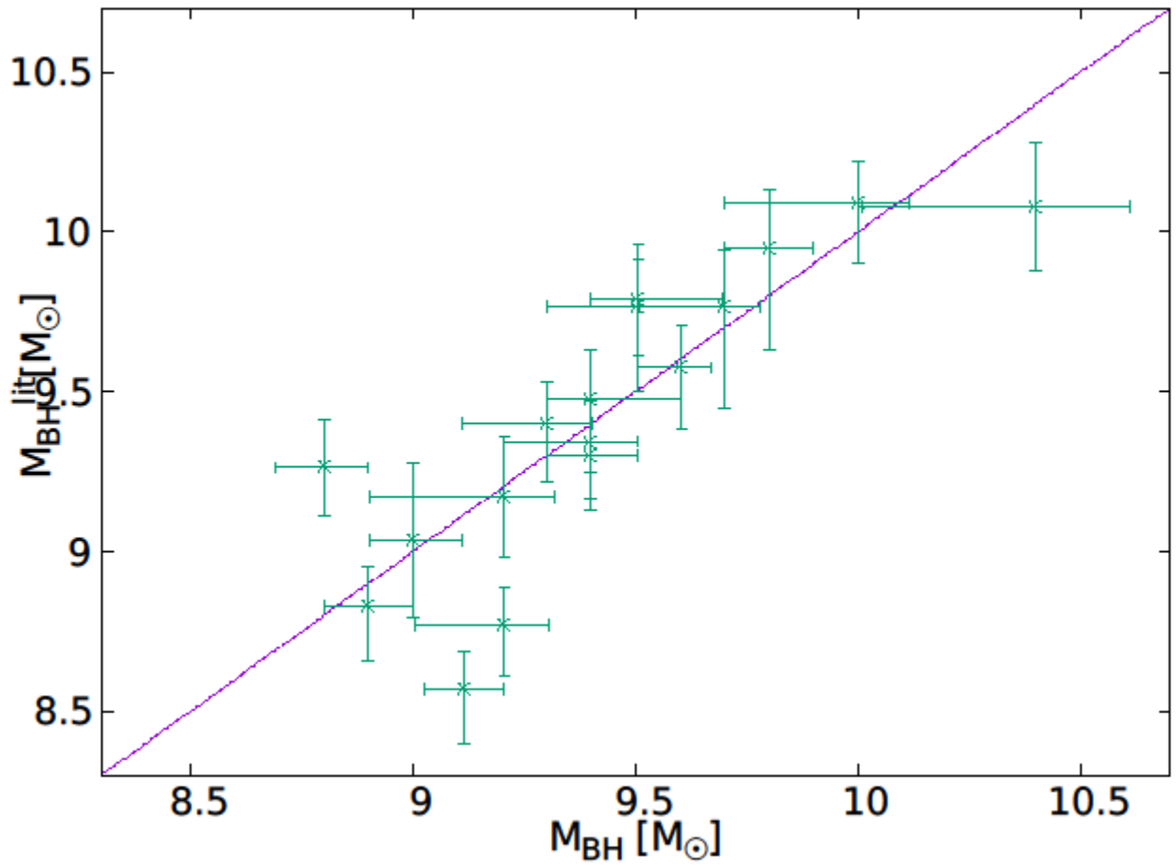
Models of continuum of weak emission-line quasars

Comparing with the Czerny et al. model

- In the main range models are consistent,
- The relativistic effect rightly suggest a difference in the ultraviolet range



To improve our method we compared WLQ with normal QSOs. For this purpose we selected sample from the Large Bright Quasar Survey (LBQS). We based on 27 quasars from Vestergaard & Osmer (2009). We chose these object because of presence of visible big blue bump



Models of continuum of weak emission-line quasars